

CHAPTER 13: Neutron Stars and Black Holes

- Type 2 Supernovae - core bounce in evolved high mass stars
 - After the bounce, a super-dense remnant remains
 - Just before the bounce, the iron atom particles (electrons and protons) in the core are crushed together
 - Neutronization
 - We're left with a ball of neutrons supported by neutron degeneracy pressure -> A neutron star!
 - Composition: Neutrons
 - Mass: 1-3 M_{sun}
 - Radius: 10 km (6 miles)
 - Density: 10^{11} kg/cm^3
- Neutron stars spin very rapidly because:
 - Angular momentum, L , is conserved
 - Mass stays the same
 - Before collapse, stars rotate once per 25 days
 - Neutron stars rotate in less than a second
- Neutron stars have very strong magnetic fields because:
 - Contracting stars carry their magnetic fields with them
 - The field intensifies with the contraction so compact objects (like neutron stars) often have strong magnetic fields
- The combination of these characteristics leads to a phenomenon called: PULSARS
 - The first PULSAR was discovered in 1967 by Jocelyn Bell, who measured this radio signal from an unresolved source... in 1974 her advisor Anthony Hewish won the nobel prize in physics for explaining it:
 - 1: Strong magnetic fields lead to hot spots at the magnetic poles
 - 2: Accelerated particles at hot spots emit beamed radiation
 - 3: If the rotation axis DOES NOT equal the magnetic field axis, then this "searchlight" rotates
 - 4: The rotating searchlight is like a lighthouse, which we see as a Pulsar if we happen to lie in the searchlight beam
 - Not all neutron stars will be observed by us as Pulsars, because the searchlight might not intercept Earth
 - The pulse period gives the rotation period of the neutron star
 - Typical periods are .03 to 1 second, or 1-30 rotations per second

$$L_{\text{before}} = L_{\text{after}}$$

$$mVr = mVr$$



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Table of Contents Astronomy Notes 2002 Update Pb 2002

1. Understanding the eBook Astronomy Notes 2002 Update Pb 2002
 - The Rise of Digital Reading Astronomy Notes 2002 Update Pb 2002
 - Advantages of eBooks Over Traditional Books
2. Identifying Astronomy Notes 2002 Update Pb 2002
 - Exploring Different Genres
 - Considering Fiction vs. Non-Fiction
 - Determining Your Reading Goals
3. Choosing the Right eBook Platform
 - Popular eBook Platforms
 - Features to Look for in an Astronomy Notes 2002 Update Pb 2002
 - User-Friendly Interface
4. Exploring eBook Recommendations from Astronomy Notes 2002 Update Pb 2002

- Personalized Recommendations
 - Astronomy Notes 2002 Update Pb 2002 User Reviews and Ratings
 - Astronomy Notes 2002 Update Pb 2002 and Bestseller Lists
5. Accessing Astronomy Notes 2002 Update Pb 2002 Free and Paid eBooks
 - Astronomy Notes 2002 Update Pb 2002 Public Domain eBooks
 - Astronomy Notes 2002 Update Pb 2002 eBook Subscription Services
 - Astronomy Notes 2002 Update Pb 2002 Budget-Friendly Options
 6. Navigating Astronomy Notes 2002 Update Pb 2002 eBook Formats
 - ePub, PDF, MOBI, and More
 - Astronomy Notes 2002 Update Pb 2002 Compatibility with Devices
 - Astronomy Notes 2002 Update Pb 2002 Enhanced eBook Features
 7. Enhancing Your Reading Experience
 - Adjustable Fonts and Text Sizes of Astronomy Notes 2002 Update Pb 2002
 - Highlighting and Note-Taking Astronomy Notes 2002 Update Pb 2002
 - Interactive Elements Astronomy Notes 2002 Update Pb 2002
 8. Staying Engaged with Astronomy Notes 2002 Update Pb 2002
 - Joining Online Reading Communities
 - Participating in Virtual Book Clubs
 - Following Authors and Publishers Astronomy Notes 2002 Update Pb 2002
 9. Balancing eBooks and Physical Books Astronomy Notes 2002 Update Pb 2002
 - Benefits of a Digital Library
 - Creating a Diverse Reading Collection Astronomy Notes 2002 Update Pb 2002
 10. Overcoming Reading Challenges
 - Dealing with Digital Eye Strain
 - Minimizing Distractions
 - Managing Screen Time
 11. Cultivating a Reading Routine Astronomy Notes 2002 Update Pb 2002
 - Setting Reading Goals Astronomy Notes 2002 Update Pb 2002
 - Carving Out Dedicated Reading Time
 12. Sourcing Reliable Information of Astronomy Notes 2002 Update Pb 2002

- Fact-Checking eBook Content of Astronomy Notes 2002 Update Pb 2002
 - Distinguishing Credible Sources
13. Promoting Lifelong Learning
- Utilizing eBooks for Skill Development
 - Exploring Educational eBooks
14. Embracing eBook Trends
- Integration of Multimedia Elements
 - Interactive and Gamified eBooks

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